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The UKIDSS-2MASS Proper Motion Survey I: Ultracool dwarfs from UKIDSS DR4

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ABSTRACT

The UKIRT Infrared Deep Sky Survey (UKIDSS) is the first of a new generation of infrared surveys. Here we combine the data from two UKIDSS components, the Large Area Survey (LAS) and the Galactic Cluster Survey (GCS), with 2MASS data to produce an infrared proper motion survey for low mass stars and brown dwarfs. In total we detect 267 low mass stars and brown dwarfs with significant proper motions. We recover all ten known single L dwarfs and the one known T dwarf above the 2MASS detection limit in our LAS survey area and identify eight additional new candidate L dwarfs. We also find one new candidate L dwarf in our GCS sample. Our sample also contains objects from eleven potential common proper motion binaries. Finally we test our proper motions and find that while the LAS objects have proper motions consistent with absolute proper motions, the GCS stars may have proper motions which are significantly under-estimated. This is due possibly to the bulk motion of some of the local astrometric reference stars used in the proper motion determination.

Key words: Astronomical data bases: Surveys – infrared: stars – Astrometry and celestial mechanics: Astrometry – Stars: low-mass, brown dwarfs – Stars: luminosity function, mass function

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1 INTRODUCTION

The study of the lowest mass stars and brown dwarfs is one of the most active areas of current Galactic research. The identification of samples of such objects allows both the low mass/substellar luminosity and mass functions to be constrained. These can in turn be used to constrain models of star and brown dwarf formation. Additionally such samples may produce interesting single or multiple objects, whose spectroscopic properties could inform atmospheric models for brown dwarfs and giant planets.

Until relatively recently the majority of proper motion surveys have concentrated on optical data. Among the first to use proper motion as a tool to select faint, nearby populations was Luyten. His work using photographic plates culminated in two large scale catalogues of high proper motion stars, the Luyten Half Arcsecond catalogue (LHS - Luyten 1979a) and the New Luyten Two-Tenths catalogue (NLTT - Luyten 1979b). Since that time much work has gone into identifying objects Luyten missed and filling in areas of poor coverage. The majority of recent wide-field proper motion surveys have used digitised photographic plate data, mostly in optical filters. Such surveys include Lepine & Shara (2008), Hambly et al. (2004), Pokorny et al. (2003) and Finch et al. (2007). Infrared proper motion surveys are an ideal instrument for producing samples of low mass stars and brown dwarfs. These objects are cool $T_{eff} < 3500K$ and hence are intrinsically faint and red. As a result they are most easily detected in the infrared. The first infrared proper motion survey was conducted by Deacon, Hambly & Cooke (2005), this used SuperCOSMOS (Hambly et al. 2001) scans of UKST *I* plates along with data from the 2 Micron All Sky Survey (2MASS - Skrutskie et al., 2006) to produce a sample of 144 low mass stars and brown dwarfs with proper motions above half an arcsecond per year. In the later paper (Deacon & Hambly 2007) the sample was expanded to more than 7000 objects with the minimum proper motion limit reduced to $0.1''/\text{yr}$. Recently Looper et al. (2008) used available overlap areas in the 2MASS data to detect two new L dwarfs using an infrared proper motion survey. Metchev et al. (2008) used 2MASS data along with those from the partially infrared Sloan Digital Sky Survey (Adelman-McCarthy et al. 2006) to identify two new T dwarfs and 22 new L dwarfs.

Recent large-scale infrared surveys such as the DEep Near Infrared Survey (DENIS - Epchtein et al., 1997)) and 2MASS took place at the end of the last decade and the start of this one. They produced CCD quality infrared data over tens of thousands of square degrees to depths of 16.5 and 15.8 respectively in the *J* band. Both surveys have led to large samples

of late type objects being identified and classified (Kirkpatrick et al. 1999, Burgasser et al. 2002, Delfosse et al. 1997). The generation of infrared surveys following DENIS and 2MASS is led by the UKIRT Infrared Deep Sky Survey (UKIDSS - Lawrence et al. 2007). This suite of surveys using WFCAM (Casali et al. 2007) on the UK Infrared Telescope (UKIRT) provides both large scale surveys for studies of galactic and extragalactic populations and deep pencil-beam surveys to examine the population of high redshift galaxies. The two surveys of the most interest to the study of low mass stars and brown dwarfs are the Large Area Survey (LAS) and the Galactic Clusters Survey (GCS). The LAS will have a final area of 4000 square degrees, will be in four passbands (Y , J , H and K) to a depth of $J=19.6$ and will also provide a second J band epoch for proper motion measurements. It covers an area away from the Galactic plane and is intended for the study of high redshift quasars and cool dwarf/subdwarfs in the field. This survey has already produced interesting results, both in terms of individual cool objects (Warren et al. 2007) and the population of low mass objects (Pinfield et al. 2008). The GCS concentrates on ten star forming regions and open clusters. In addition to the four colours used in the LAS, Z band data is available and a second K epoch will provide proper motion data for the clusters. So far studies such as Lodieu et al. (2006) and Lodieu et al. (2007) have used the data released to produce mass functions for Upper Sco and the Pleiades respectively. The GCS can be used also for the study of objects lying in the foreground of the target clusters (Lodieu et al., in preparation). The UKIDSS data is released through the WFCAM Science Archive (WSA - Hambly et al. 2008) and is available to all European Southern Observatories (ESO) members. The current release is Data Release 4 (DR4). Additionally astronomers from non-ESO countries can access the data after an 18 month delay. The soon to be operational Visual and Infrared Survey Telescope for Astronomy (VISTA - Emerson et al. 2004) will provide an even faster infrared surveying capability than WFCAM. It will cover the whole southern hemisphere in two bands (J and K) including several thousand square degrees with additional Y and/or H band photometry.

2 METHOD

Our method for selecting candidate ultracool dwarfs is almost identical to that used in Deacon et al. (2005) and Deacon & Hambly (2007). The first selection process is to identify UKIDSS objects with counterparts in 2MASS within a distance of roughly $4-5\sigma$ (where σ

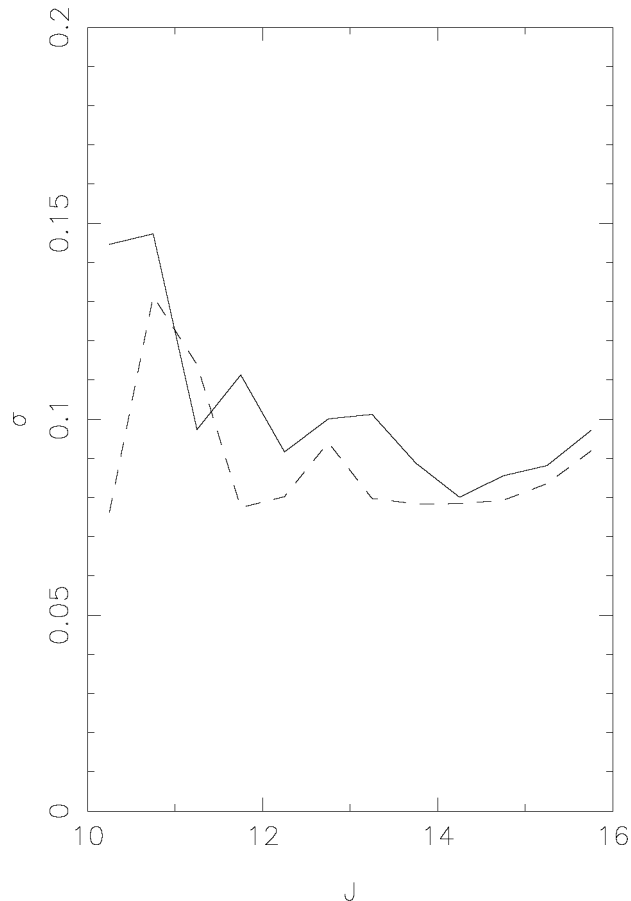


Figure 1. The astrometric errors between the UKIDSS and 2MASS data. The solid line represents the error in Right Ascension and the dashed line the error in Declination.

is the positional error between the surveys). Clearly to be able to carry out this selection we must first identify the typical positional errors between the surveys. This was done by comparing the positions of objects between the two surveys and calculating the standard deviation between the two while removing non-Gaussian outliers. A plot of the errors vs. J band magnitude is shown in Figure 1. It is reasonable to assume that above the saturation limit ($J=10.5$) the positional errors are of the order of 0.1 arcseconds in each axis. Hence a minimum positional shift of 0.6 arcseconds will have a significance of $\geq 5\sigma$.

In addition to the astrometric cut, only objects with good stellar profiles (i.e. a classification statistic between -3 and 3) were selected. We also excluded objects with magnitudes brighter than the saturation limit for each filter and those more than one magnitude fainter than the 2MASS detection limit of $J=15.8$. Next we examined the $(Y - J)$ vs. $(J - H)$ colour-colour plot found in Leggett et al. (2005) and set a cut that either $(Y - J)$ had to be greater than 0.7 or $(Y - J)$ had to be greater than 0.3 and $(J - H)$ less than zero. This allowed us to exclude the main mass of main sequence stars while still covering the areas

where L and T dwarfs (and the Baraffe et al. 2003 models for Y dwarfs) are expected to lie. All of these cuts were enacted using the SQL query shown in Appendix A. This was submitted to the WFCAM Science Archive which returned our unpaired UKIDSS sample.

Next we attempted to identify 2MASS companions to the unpaired UKIDSS objects. This was done by searching through the appropriate 2MASS files and identifying any good stellar object within 100 arcseconds of the UKIDSS object with a similar J band magnitude (within one magnitude). To ensure that each 2MASS object was not a non-moving coincidence background object, we then checked to ensure each object did not have a UKIDSS companion within the 0.6 arcsecond movement cut. This was done using the WSA CrossID function. Any 2MASS object with a UKIDSS pair within 0.6 arcseconds was excluded from the sample. Next the UKIDSS images of all the candidates were inspected by eye to remove any objects which could have poor astrometry resulting from deblended images or nearby bright stars.

To ensure accurate proper motion estimates we attempted to calculate individual astrometric solutions for each object in our sample. To do this we used the WSA Cross-ID function to identify all good stellar images within five arcminutes of each object in our sample. These objects provide a set of astrometric reference stars for each star in our sample. We then cross referenced these reference stars with the 2MASS catalogue so we had their positions in both datasets. For each star in our sample we selected reference stars with J band brightnesses within one magnitude of the target and used these to do a standard six parameter plate-to-plate fit. This allows us to correct for any offsets between the astrometric systems and calculate local positional errors. If a target had too few reference stars (five or fewer) associated with it, no plate-to-plate fit was carried out and a magnitude dependent error estimate was taken from the global error estimate shown in Figure 1. Whether an object’s astrometry uses the local or global astrometric solutions is indicated in the data tables (see Appendix B). Only objects with proper motions more significant than 5σ were included in the final sample. However an astrometric solution calculated using too few reference stars will underestimate the positional errors. Hence we carried out a series of simulations to estimate difference between the positional errors found using a set of n_{ref} ($\sigma_{measured}$) and the true positional error (σ_{true}) which we set as an input to our simulations. It was found that the following relationship was a good estimate for the underestimation of the positional errors,

$$\frac{\sigma_{true}}{\sigma_{measured}} \approx 1 + \frac{19.8}{n_{ref}^{1.5}} \quad (1)$$

Hence this equation was used to apply a correction factor to all the astrometric solutions calculated using sets of reference stars. The main aim of this study is to examine the population of ultracool dwarfs in the field. This is complicated by the fact that in the GCS areas, our 5σ lower proper motion cut could include some objects which are members of the target clusters. Hence for the GCS data we have set a minimum proper motion of 80 milliarcseconds per year. This excludes objects with proper motions similar to all the GCS clusters with the exception of the Hyades (which does not appear in the dataset we use).

Our survey covers all the area included in the UKIDSS DR4 for the LAS and the GCS. These surveys consist of 993 sq.deg. and 148 sq.deg. of sky respectively.

3 RESULTS

The Large Area Survey (LAS) sample produced 213 candidate objects of which 44 are previously identified by other surveys. The Galactic Clusters Survey (GCS) produced a sample of 54 objects, 12 of which are previously known. A list of all the objects identified can be found in Appendix B. Figure 2 shows infrared colour-colour diagrams for both survey samples. In both cases it is clear that there is only one object which is both redder than $Y - J = 1.1$ and bluer than $J - H = 0.5$. Comparing these with colours from Hewett et al. (2006) we can conclude that there is only one object with a spectral type later than T3 in our sample (our full sample includes all objects found in the LAS and GCS). One clear difference between the two samples of objects found in the LAS and those found in the GCS is the appearance of several objects above the main M dwarf locus in the GCS sample. By comparison the LAS sample has a tight stellar locus which becomes less dense towards the L dwarf regime (redder than $Y - J = 1.1$). One member of this second locus is NLTT 42735, an object which appears in the initial DENIS low mass stars sample of Crifo et al. (2005). They use spectroscopy to identify NLTT 42735 (UGCS2MASS1625–2400) as a reddened F–K type star and hence exclude it from their final sample. They point out that its reddening is probably due to its proximity to a known molecular cloud. Checking our seven objects (UGCS2MASS0408+2447, UGCS2MASS0536-0454, UGCS2MASS0433+2933, UGCS2MASS0541–0305, UGCS2MASS1625–2400, UGCS2MASS1631–2404 and UGCS2MASS1637–2200) which lie above the main stellar locus using the SIMBAD database reveals that five are within ten arcminutes of known molecular clouds or dark nebulae. Given that this second locus only appears in the GCS (which samples clusters which lie close to star formation

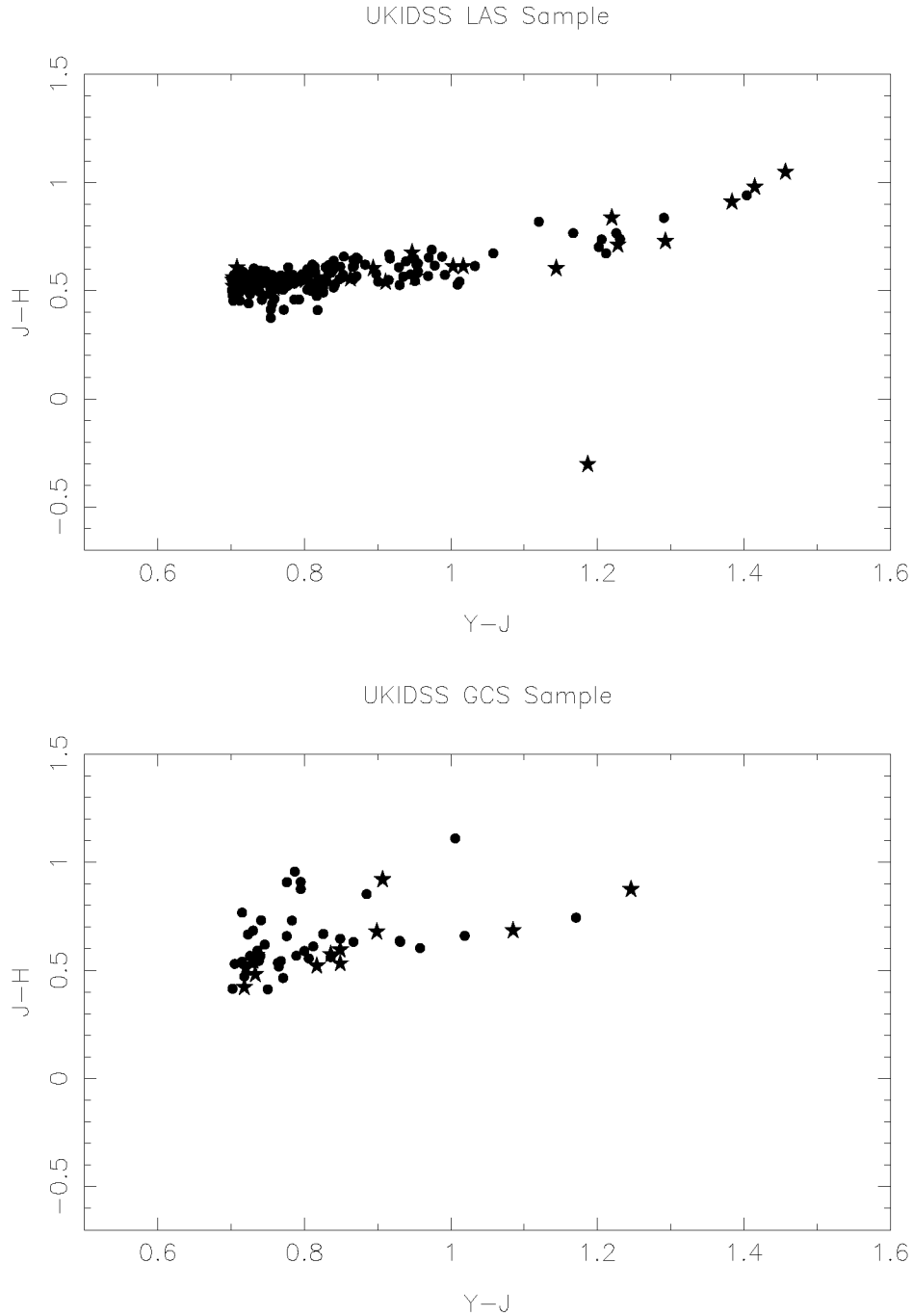


Figure 2. The UKIDSS colour-colour diagrams for our two samples. Dots represent newly discovered objects and stars those found in previous surveys. Note the second locus in the GCS data which lies above the main locus. We believe this consists of highly reddened early type stars behind molecular clouds.

regions) and not in the LAS (which samples an area of high Galactic latitude) we believe we can assume that it consists of earlier-type stars which lie behind the star formation regions and are reddened by the gas and dust. As we do not believe these objects are true low mass stars or brown dwarfs, we recommend that they are excluded from any cool dwarf sample derived from our data. As a number of other objects appear to lie only slightly above the

Table 1. A list of all objects in our sample redder than $Y - J=1.1$. The astrometric solutions were calculated using a fit to local reference stars. All photometry uses the standard UKIDSS filters (Hewett et al. 2006). Citation key — a:Schneider et al. (2002), b: Fan et al. (2000), c: Bouy et al. (2003), d: Reid et al. (2008), e: Burgasser et al. (2004), f: Wilson et al. (2003), g: Kirkpatrick et al. (2000), h: Knapp et al. (2004), i: Hawley et al. (2002).

Name	Position	μ "/yr	P.A. °	σ_μ "/yr	σ_{PA}	Y	J	H	K	note
ULAS2MASS0001+1535	00 01 12.24 +15 35 34.3	0.218	144.0	0.029 ¹	8.4	16.877	15.462	14.483	13.623	h
ULAS2MASS0054-0031	00 54 06.66 -00 31 03.2	0.248	129.8	0.029 ¹	6.9	16.857	15.629	14.917	14.302	a
ULAS2MASS0205+1251	02 05 03.66 +12 51 42.0	0.362	93.8	0.025 ¹	4.3	17.018	15.561	14.513	13.637	g
ULAS2MASS0207+1355	02 07 35.73 +13 55 54.9	0.316	123.3	0.017 ¹	3.6	16.589	15.369	14.532	13.829	i
ULAS2MASS0219+0506	02 19 22.05 +05 06 30.8	0.184	84.1	0.013 ¹	5.5	16.022	14.855	14.089	13.453	
ULAS2MASS0330-0025	03 30 35.32 -00 25 37.4	0.509	131.1	0.015 ¹	1.7	16.508	15.215	14.487	13.771	b
ULAS2MASS0344+0111	03 44 08.88 +01 11 23.9	0.209	211.1	0.030 ¹	6.1	15.698	14.554	13.952	13.446	c
UGCS2MASS0409+2104	04 09 09.57 +21 04 37.9	0.184	150.6	0.007 ¹	2.4	16.622	15.376	14.502	13.781	i
ULAS2MASS0835+0548	08 35 58.24 +05 48 30.7	0.110	257.6	0.009 ¹	7.7	15.705	14.499	13.762	13.127	d
ULAS2MASS0843+1024	08 43 33.31 +10 24 43.0	0.593	165.3	0.011 ¹	1.5	15.989	14.777	14.105	13.551	
ULAS2MASS1211+0406	12 11 30.11 +04 06 08.1	0.225	195.7	0.035 ¹	6.5	16.808	15.517	14.680	13.949	
ULAS2MASS1231+0847	12 31 46.99 +08 47 25.8	1.590	228.1	0.023 ¹	0.9	16.340	15.153	15.456	15.552	e
ULAS2MASS1308+0818	13 08 30.97 +08 18 52.5	0.233	281.1	0.022 ¹	3.4	16.312	15.192	14.373	13.792	
ULAS2MASS1346+0842	13 46 07.37 +08 42 34.1	0.235	246.6	0.042 ¹	8.9	16.744	15.518	14.752	14.115	
ULAS2MASS1407+1241	14 07 53.45 +12 41 10.4	0.337	280.5	0.022 ¹	3.3	16.737	15.333	14.392	13.632	d
ULAS2MASS1422+0827	14 22 57.10 +08 27 50.4	0.592	195.6	0.012 ¹	1.5	16.240	15.009	14.271	13.609	
ULAS2MASS1448+1031	14 48 25.75 +10 31 58.1	0.278	117.5	0.018 ¹	5.1	15.804	14.420	13.510	12.674	f
ULAS2MASS1452+1114	14 52 01.96 +11 14 56.9	0.394	139.4	0.027 ¹	4.1	16.771	15.569	14.867	14.280	
UGCS2MASS1630-2120	16 30 17.69 -21 20 01.4	0.153	259.8	0.013 ¹	3.8	15.695	14.524	13.780	13.174	

Table 2. Photometry and astrometry for the two objects ULAS2MASS1253+0740a and ULAS2MASS1253+0740b which share a common proper motion. The astrometric solutions were calculated using fits to local reference stars.

Name	Position	μ "/yr	P.A. °	σ_μ "/yr	σ_{PA}	Y	J	H	K
ULAS2MASS1253+0740a	12 53 49.36 +07 40 04.5	0.162	269.6	0.026 ¹	5.6	14.559	13.807	13.323	12.870
ULAS2MASS1253+0740b	12 53 49.69 +07 40 00.6	0.156	269.5	0.029 ¹	7.6	14.977	14.168	13.669	13.181

main stellar locus we would advise anyone using our GCS sample to bear in mind that these objects may be reddened early type stars also.

3.1 Ultracool Dwarfs

The prime motivator for this survey is to identify a clean sample ultracool dwarfs. To further this aim we selected objects in our sample which were redder than $Y - J=1.1$. According to Hewett et al. (2006), this should give us a sample completely free from M dwarfs and earlier spectral types. The objects identified are listed in Table 1 along with their astrometry and photometry. Of the total of nineteen objects listed, eleven are previously known (one of which is a previously identified T dwarf). Hence of the eighteen objects with L dwarf-like colours only ten are previously known. Clearly spectroscopic identification of the previously unidentified eight objects is required, but even if only a few are genuine L dwarfs, this would indicate incompleteness in our current knowledge of even the relatively bright L dwarf population.

Table 3. Photometry and astrometry for objects in our sample which have a common proper motion companion found in another study. Astrometric solutions were calculated using fits to local reference stars. Some companion objects were brighter than the UKIDSS saturation limits. For these the photometry is drawn from 2MASS and they can be recognised by their lack of a Y magnitude. Citation key - a: Tycho catalogue (Hog et al. 1998), b: Argelander (1903), c: Lepine & Shara (2005), d: Hipparcos catalogue (Perryman 1997), e: Luyten (1979b), f: Giclas, Slaughter & Thomas (1959), g: Hawley et al. (2002).

Name	Position	μ "/yr	P.A. °	σ_μ "/yr	σ_{PA}	Y	J	H	K	note
ULAS2MASS0041+1341 NLTT 2274	00 41 54.44 +13 41 34.1 00 41 55.45 +13 41 16.4	0.254 0.268 ^c	230.6 228.5 ^c	0.015 0.008 ^c	15.389 1.7 ^c	14.372	13.760 10.164	13.208 9.574	g 9.347	b
ULAS2MASS0158-0025 BD-01 266	01 58 05.95 -00 25 42.8 01 58 07.36 -00 25 10.8 ^a	0.132 0.141 ^a	170.4 172.9 ^a	0.021 ¹ 0.002 ^a	7.9 1.1 ^a	15.282	14.575 8.096	14.063 7.632	13.636 7.478	b
ULAS2MASS0207+1355 G 73-26	02 07 35.73 +13 55 54.9 02 07 37.48 +13 54 49.5	0.316 0.321 ^c	123.3 125.4 ^c	0.017 ¹ 0.008 ^c	3.6 1.4 ^c	16.589	15.369 9.196	14.532 8.569	13.829 8.307	g f
ULAS2MASS0946+1116 NLTT 22538	09 46 12.09 +11 16 31.1 09 46 12.64 +11 16 37.1	0.164 0.183 ^c	265.9 263.0 ^c	0.027 ¹ 0.008 ^c	10.8 2.5 ^c	16.407	15.677 10.233	15.090 9.698	14.698 9.554	
ULAS2MASS1159+0706 NLTT 29211	11 59 48.15 +07 06 59.1 11 59 48.58 +07 07 08.6	0.201 0.204 ^c	301.1 302.4 ^c	0.023 ¹ 0.008 ^c	6.2 2.3 ^c	16.027	15.304 17.526	14.801 17.575	14.399 17.491	e
ULAS2MASS1259+0651 LSPM J1259+0651	12 59 37.59 +06 51 18.9 12 59 39.34 +06 51 25.6 ^c	0.444 0.483 ^c	240.3 240.7 ^c	0.032 ¹ 0.008 ^c	5.0 0.9 ^c	15.045	14.289 13.875	13.851 13.461	13.422 13.061	c
ULAS2MASS1320+0957 NLTT 33793	13 20 41.49 +09 57 49.7 13 20 50.13 +09 55 58.3 ^d	0.256 0.288 ^d	238.7 240.0 ^d	0.037 ¹ 0.002 ^d	6.8 0.3 ^d	14.484	13.654 7.889	13.111 7.323	12.627 7.218	e
ULAS2MASS1327+0916 LSPM J1327+0916	13 27 26.77 +09 16 05.6 13 27 28.51 +09 16 32.4 ^c	0.146 0.161 ^c	238.3 244.2 ^c	0.012 ¹ 0.008 ^c	5.6 2.8 ^c	15.300	14.539 10.68	13.994 10.07	13.578 9.83	c
UGCS2MASS0336+2233 NLTT 11342	03 36 25.95 +22 33 17.2 03 36 23.63 +22 33 27.3 ^c	0.243 0.244 ^c	132.7 134.5 ^c	0.007 ¹ 0.008 ^c	1.7 1.9 ^c	15.808	14.991 11.581	14.471 11.019	13.977 10.762	e
UGCS2MASS0342+2248 HD 22992	03 42 10.17 +22 48 44.6 03 42 15.92 +22 47 11.2 ^d	0.089 0.079 ^d	166.9 164.1	0.015 ¹ 0.001 ^d	10.7 0.9 ^d	16.355	15.543 6.512	14.932 6.343	14.469 6.313	d

3.2 Common Proper Motion Objects

We utilised the same technique that was used in Deacon & Hambly (2007) to search for objects in our sample that may have a common proper motion companion within our sample. This means we searched for objects separated by less than three arcminutes with proper motions within two sigma of each other. This produced one potential pair, ULAS2MASS1253+0740a and ULAS2MASS1253+0740b. These are separated by only 6.3 arcseconds (only just outside the 2MASS proximity limit for good quality photometry and astrometry). The details of this pair can be seen in Table 2.

Additionally we checked all of our objects against the Simbad database to identify any from previous surveys which may be common proper motion companions. Using the same criteria as before we identified ten potential common proper motion pairs the details of which can be found in Table 3. Of these two have measured Hipparcos parallaxes, NLTT 33793 (a K8 star) with $\pi=26.22\pm1.68$ milli-arcseconds and HD 22992 (an F2 star) with $\pi=18.79\pm1.05$ milli-arcseconds. At these parallaxes, (with angular separations of 2.83 and 2.05 arcminutes respectively) both systems would have separations of roughly 6500AU. However in Section 3.4 we calculate distance estimates for these two objects. These lead us to conclude that HD

Table 4. The astrometry and photometry for the two objects for which spectra were taken. For both, the first line includes the astrometry and UKIDSS photometry and the second line the 2MASS photometry. No relative local astrometry was possible in the fields of these stars owing to a lack of suitable reference stars, so the global astrometry was used for proper motion measurements. Citation key - a: Luyten NLTT, b: Lepine & Shara (2005)

Name	Position	μ "/yr	P.A. °	σ_μ "/yr	σ_{PA}	<i>Y</i>	<i>J</i>	<i>H</i>	<i>K</i>	note
ULAS2MASS1017+0711	10 17 23.35 +07 11 03.9	0.159	158.0	0.024 ¹	6.7	15.396 12.504	14.656 11.834	14.103 11.488	13.695	a
ULAS2MASS1354+0846	13 54 08.67 +08 46 08.7	0.230	283.7	0.014 ²	3.2	12.976 12.193	12.146 11.605	11.587 11.156	11.144	b

229920 and UGCS2MASS0342+2248 are not physically related but that NLTT 33793 and ULAS2MASS1320+0957 probably are.

3.3 Spectroscopic Follow-up

The sample presented here uses data from the UKIDSS Data Release 4. However before these data were available we used the same method on the smaller Data Release 2 to identify potentially interesting objects. As a result we applied for service observations on the UK Infrared Telescope (UKIRT) for a number of objects whose colours and magnitudes suggested they were within 30pc. Of these, two objects had spectra taken by UKIRT staff on the nights of the ninth and tenth of January 2008. These spectra were obtained using the IJ grating on the UIST spectrograph and consisted of four individual spectra jittered in an ABBA pattern. Standard star spectra and calibration frames were also taken. The data were reduced using the Figaro Starlink package. Figures 3 and 4 show the objects ULAS2MASS1017+0719 (previously known as NLTT 23911) and ULAS2MASS1354+0846 (previously known as LSPM J1354+0846) respectively. Table 4 shows the UKIDSS photometry and our calculated astrometry for the two objects. We used the equivalent widths of K I lines at 1.169 and 1.179 microns and the values for those lines quoted in Cushing et al. (2005) for each spectral type to determine spectral classification. With an EW (in angströms) of 6.9 for the 1.179 μ m line and 2.4 for the 1.169 μ m line we deduce that ULAS2MASS1017+0719 is of spectral type M6-8, most probably M7. With equivalent widths of 7.1 and 3.5 respectively we estimate that ULAS2MASS1354+0846 is of the same spectral type. If we compare the 2MASS photometry of our objects with that of the M7 dwarf SO0253+1652 (which Henry et al., 2006 found to lie at $3.85 \pm_{0.05}^{0.03}$ pc) we can produce a rough distance estimate for both objects. SO0253+1652 has 2MASS magnitudes $J=8.394$, $H=7.883$ and $K_s=7.585$. Based on this photometry we estimate that ULAS2MASS1017+0719 lies at roughly 23.7pc and

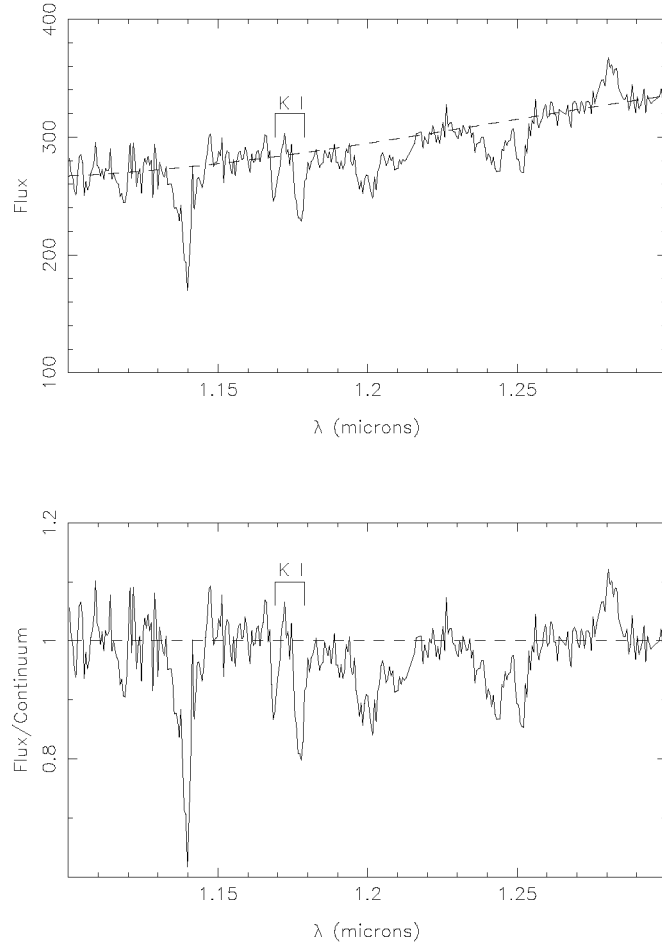


Figure 3. The infrared spectrum of ULAS2MASS1017+0719 also known as NLTT 23911. The top panel shows the flux calibrated spectrum and the lower panel the continuum divided spectrum. The spectral lines used for the classification are marked K I

ULAS2MASS1354+0846 at 21.0pc. Clearly these are rough estimates and are inferior to trigonometric parallax measurements.

3.4 Photometric Distance Estimates

While a list of potential nearby cool dwarfs is useful, photometric distance estimates can help to identify the nearest objects in our sample. To do this we took the data from Hewett et al. (2006) and Leggett (private communication) for the WFCAM colours of cool dwarfs and combined these data with those from Golimowski et al. (2004) (which is itself based on a series of trigonometric parallax measurements from the literature) to produce a series of relations. The first set of these relate colours to spectral type. These are plotted in the lower panels of Figure 5. Separate polynomials for the M and L and T dwarf regimes are fitted, these are shown in Table 5. The $Y - J$, $J - H$ and $H - K$ colours were used to calculate a separate spectral type for each for the target object. As many of these relations have no unique solution for some colours, more than one spectral type for each colour was

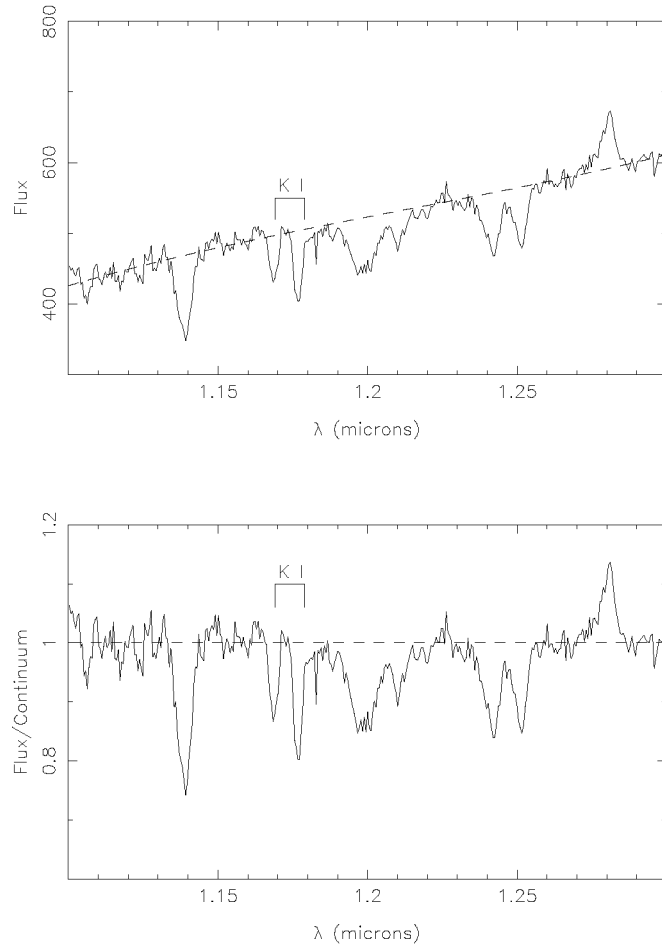


Figure 4. The infrared spectrum of ULAS2MASS1354+0846 also known as LSPM J1354+0846. The top panel shows the flux calibrated spectrum and the lower panel the continuum divided spectrum. The spectral lines used for the classification are marked K I

normally calculated. Next the spectral types were grouped together so that three spectral types from the three colours were in the close agreement. In some cases there was more than one valid grouping of estimated spectral types. Each group was then averaged to calculate the estimated spectral type. If there is more than one calculated spectral type, other data must be used to select the correct estimation. Here we use i and z photometry from the Sloan Digital Sky Survey. The fourth colour relation we use plots $i - z$ vs. spectral type. We use this relation along with Sloan photometry and choose the calculated spectral type where the photometry and the colour relation are in best agreement. Once this spectral type was calculated the observed magnitudes in each passband could be compared with the relation between spectral type and the absolute magnitude in that particular passband (see Table 5 and the top 4 panels of Figure 5). A distance from each passband was then calculated and these were averaged to produce a final distance estimate.

Clearly the relations must be tested to ensure they are reliable. To do this we first took the measured spectral types for our potential L and T dwarfs (see Table 1). We then com-

Table 5. The colour/magnitude to effective temperature relations used in this work. They are of the form $X = \sum_i a_i SpT^i$ where M0 is spectral type 0, L0 spectral type 10 and T0 spectral type 20.

Relation	a_0	a_1	a_2	a_3	a_4	a_5	a_6
M_Y	4.800e+00	2.147e+00	-2.844-01	2.035e-02	-4.502e-04	-1.167e-05	4.384e-07
M_J	4.074e+00	2.394e+00	-3.506e-01	2.488e-02	-5.166e-04	-1.447e-05	5.090e-07
M_H	3.492e+00	2.345e+00	-3.291e-01	2.233e-02	-4.442e-04	-1.304e-05	4.513e-07
M_K	3.276e+00	2.306e+00	-3.243e-01	2.154e-02	-4.130e-04	-1.261e-05	4.277e-07
$Y - J$	7.255e-01	-2.470e-01	6.613e-02	-4.531e-03	6.634e-05	2.802e-06	-7.063e-08
$J - H$	5.81e-01	4.888e-02	-2.151e-02	2.543e-03	-7.233e-05	-1.440e-06	5.768e-08
$H - K$	2.167e-01	3.950e-02	-4.808e-03	7.891e-04	-3.127e-05	-4.204e-07	2.358e-08
$i - z$	2.037e-01	2.891e-01	-1.951e-02	7.997e-04	2.458e-06	-6.175e-07	5.917e-09

pared these with each object's spectral calculated by our method. We found that in general we had a standard deviation of roughly 1.5 subtypes. We also used objects common between our sample and Reid et al. (2008) to estimate the accuracy of our distance estimates. These had calculated photometric distances from 2MASS data and from Cruz et al. (2003). We compared these distances with our own and found that (once the errors on the Reid et al. distance calculations had been taken into account) our calculated distances were accurate to roughly 20%. We recommend that these relations are only used for initial sample selection. We do not believe these are accurate enough to quote values for distances. Table 6 shows a list of all the objects in our sample which we believe to be within 30pc based on our photometric distance relations. Our sample includes both the objects which we spectroscopically classified. In the previous section we calculated the distances to these objects using an independent method, both were estimated to be M7 dwarfs at roughly 20pc. Here, both have distance estimates of roughly 20pc with one estimated to be an M7 dwarf and one an M6.5 dwarf. In Section 3.2 we identified two objects which shared a common proper motion with nearby stars with Hipparcos parallaxes. Using the photometric relations we find that UGCS2MASS0342+2248 lies at roughly 80.2 parsecs and ULAS2MASS1320+0957 at approximately 38.2 parsecs. As HD229920 - the common proper motion companion of UGCS2MASS0342+2248 - lies at 53pc, we must conclude these objects are unlikely to be related. However NLTT 33793 lies at 38.2pc, a good match to our distance estimate for ULAS2MASS1320+0957.

4 DISCUSSION

To provide a simple estimate of completeness we plotted a cumulative histogram of objects' proper motions. Assuming the stellar velocity and positional distributions do not change

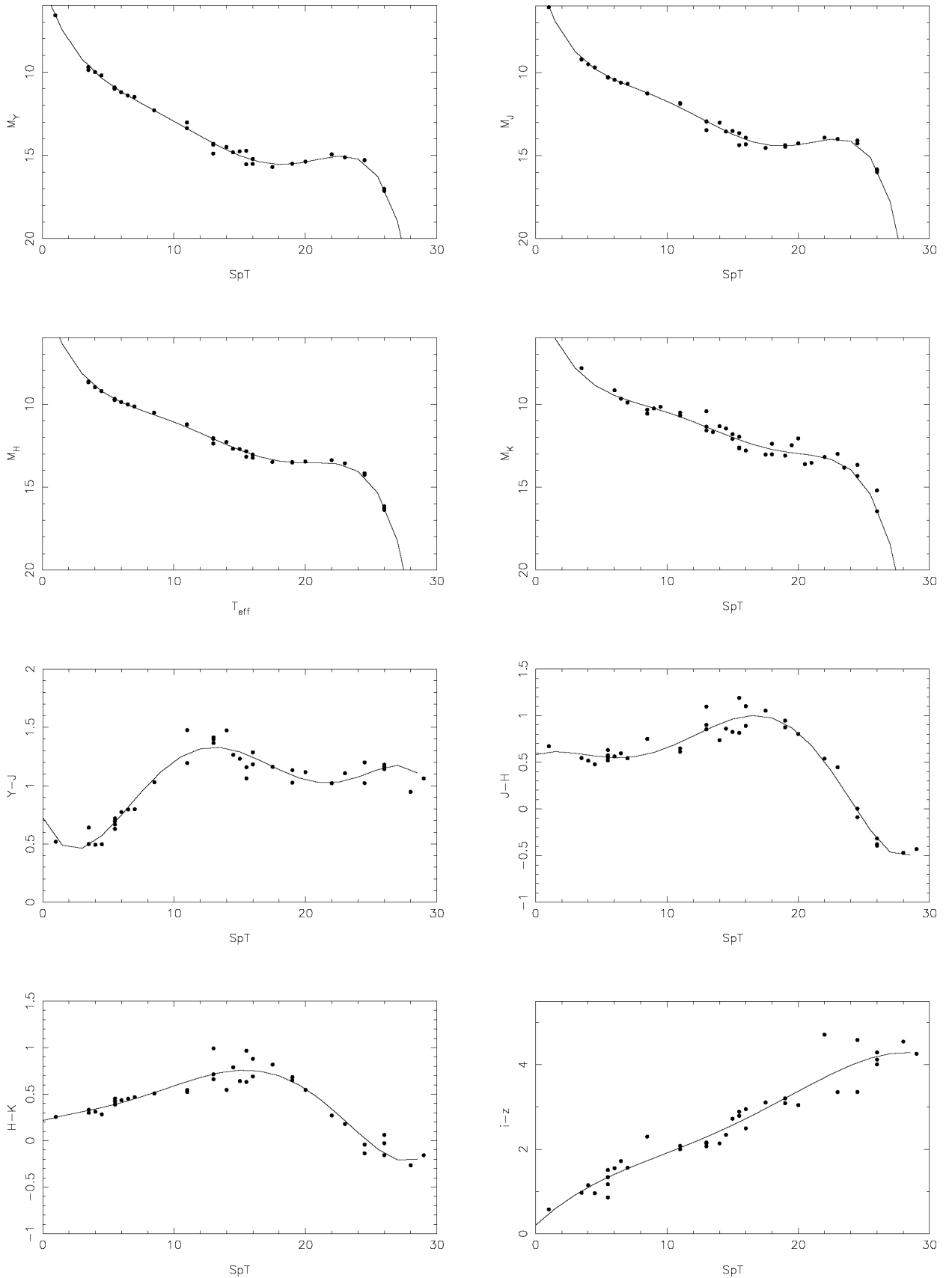


Figure 5. Plots of the data and fits for the colour magnitude relations. The data here includes the colours from Hewett et al. (2006) and absolute magnitudes from Golimowski et al. (2004). M0 is spectral type 0, L0 spectral type 10 and T0 spectral type 20.

Table 6. Objects with photometric distance estimates within 30pc according to our photometric distance relations. The first spectral type and distance are calculated using our distance estimates, the second spectral type and distance estimate come from another source. Citation key - a: Burgasser et al. (2004), b: Burgasser et al. (2006), c: Wilson et al. (2003), d: Reid et al. (2008), e: Tinney (1993), f: Tinney (1995), g: Lepine & Shara (2005), h: this work, i: Hawley et al. (2002), j: Luyten (1979b), k: Kirkpatrick et al. (1997), l: Cruz et al. (2003), m: Knapp et al. (2004), n: Phan-Bao & Bessell (2006), o: Kirkpatrick et al. (2000).

Name	μ "/yr	<i>Y</i>	<i>J</i>	<i>H</i>	<i>K</i>	SpT	d (pc)	SpT	d (pc)	note
ULAS2MASS1231+0847	1.590	16.350	15.166	15.451	15.544	T5.5	9.1	T5.5		a,b
ULAS2MASS1448+1031	0.278	15.813	14.425	13.520	12.695	L4	15.9	L4	19.6 \pm 4.0	c,d
ULAS2MASS1510-0241	0.399	13.486	12.534	11.974	11.313	M9	16.7	M9	16.34 \pm 1.25	e,d,f
ULAS2MASS0002+0115	0.455	12.894	12.093	11.545	11.074	M6.5	19.3	M6.5	21.0 \pm 2.52	j,n
ULAS2MASS0001+1535	0.218	16.877	15.462	14.483	13.623	L5	20.8	L4		m
ULAS2MASS1354+0846	0.230	12.979	12.151	11.600	11.153	M7	21.6	M7		g,h
ULAS2MASS1134+0022	0.511	13.689	12.790	12.181	11.701	M8	21.7	M9	18.9 \pm 1.2	i,d
ULAS2MASS1017+0719	0.188	13.213	12.407	11.872	11.424	M6.5	22.6	M7		j,h
ULAS2MASS0205+1251	0.362	17.018	15.561	14.513	13.637	L5	22.8	L5	26.6 \pm 3.0	o,l
UGCS2MASS0339+2457	0.178	13.608	12.772	12.223	11.732	M7	23.5	M8		k
ULAS2MASS1407+1241	0.337	16.712	15.319	14.381	13.617	L4.5	23.6	L1	47.4 \pm 12.5	d
UGCS2MASS0409+2104	0.184	16.622	15.376	14.502	13.781	L4	24.8	L3		o
UGCS2MASS0354+2316	0.206	13.903	13.054	12.469	11.982	M7.5	25.5			g
ULAS2MASS1457+1102	0.117	13.169	12.466	12.013	11.616	M6	26.5			
ULAS2MASS1546+0317	0.113	13.965	13.155	12.582	12.114	M7	27.7			
UGCS2MASS0345+2540	0.109	15.009	13.922	13.246	12.663	M9.5	28.4	L0	26.95 \pm 0.36	k,d

significantly, the cumulative number of objects should scale as μ^{-3} (see Figure 6). A line showing this scaling is also plotted. It is clear that the completeness begins to drop off below about $\mu = 0.2''/\text{yr}$. This will be due to distant, fainter objects being excluded from our sample due to magnitude limits. At the lowest proper motions higher astrometric errors may also cause some objects to be excluded. To check our completeness for cool objects we used the Dwarf Archives website to identify any potential L and T dwarfs in our survey area with brightnesses within our detection range. Eleven L dwarfs of sufficient brightness were identified in our survey area. Of these one (GD 165B) did not appear in our sample. The reason for this is that the 2MASS proximity flag for this object was below six arcseconds due to the closeness of its binary companion. Of the known T dwarfs only 2MASS12314753+0847331 lies in the survey area, this object is detected. Hence we detect all the known, single L and T dwarfs brighter than the 2MASS limit in our survey area.

In order to examine the accuracy of our proper motions we cross-referenced our catalogue with that of Lepine & Shara (2005). In the LAS sample there were fifteen common objects between the two surveys. Our proper motions of these agreed well with those of Lepine and Shara. However there were large offsets found in two of the four objects in both Lepine and Shara and our GCS sample. The reason for this is that we are measuring relative proper motions while Lepine & Shara measure absolute proper motions. Put simply both our relative astrometry and the UKIDSS astrometry use stars to define our reference frame, Lepine and

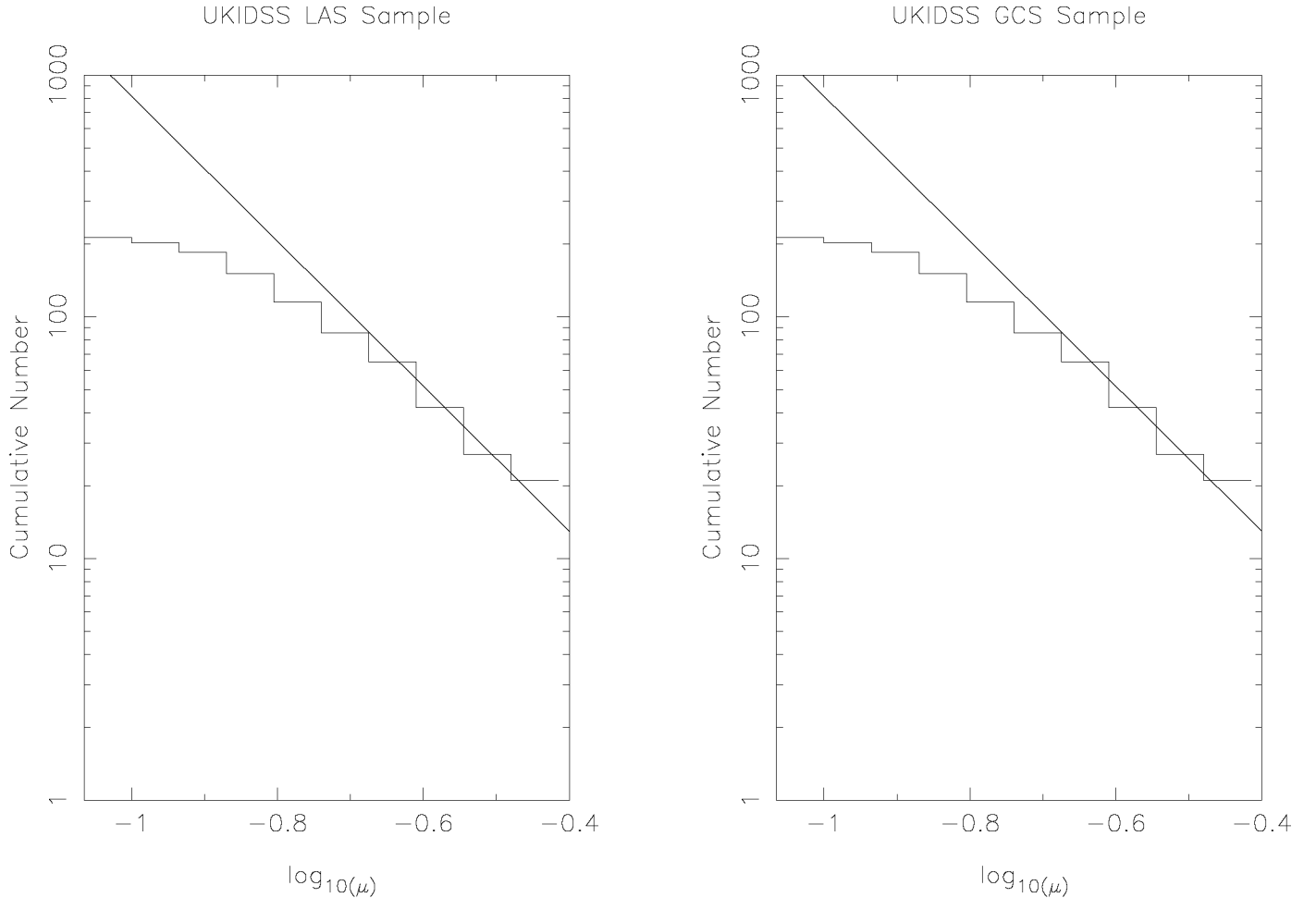


Figure 6. Cumulative histograms for both samples. With a survey with no incompleteness the number should scale as μ^{-3} (indicated by the solid line). It is clear that below $0.2''/\text{yr}$ our surveys deviates from this trend and is hence incomplete.

Shara use quasars to set their astrometric reference frame. This leaves us vulnerable to the proper motions of our reference stars introducing an error to our astrometric system. As the Galactic Cluster Survey is taken in the area of known open clusters, there will be a distinct bulk motion of a large number of potential reference stars. Hence there will be an offset in our proper motions measurements.

The objects found in the UKIDSS GCS are probably unrelated to the target clusters. As we have chosen a minimum proper motion of 80 milliarcseconds per year all our objects have proper motions at least $2\text{--}3\sigma$ higher than the proper motions of nine out of the ten target clusters and star forming regions. The one exception is the Hyades, however the core area of the Hyades does not yet have the UKIDSS GCS multiband observations required for this survey. Additionally several objects are listed as being cluster members in Table B2 (this cluster membership flagging comes from their descriptions in SIMBAD). These objects

are listed as either members of the Pleiades or α Per. If our proper motion measurements for these objects are correct it would indicate that they are not members of these clusters.

5 CONCLUSIONS

Our attempt to utilise the currently available UKIDSS data for a proper motion survey has successfully detected all the expected, previously known single L and T dwarfs in our LAS survey area. Additionally we have found seven objects in the LAS and one in the GCS that are good candidate L dwarfs. Finally, initial spectral follow-up of our sample has classified two previously known objects. Rough distance estimates put these as being within 25pc. In further iterations of this work with future UKIDSS data releases we will attempt to tie our proper motions to absolute proper motions, avoiding the offsets in the Galactic Clusters Survey sample. This work could easily be extended to cover the upcoming surveys by the VISTA telescope.

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APPENDIX A: SQL QUERY USED

```
SELECT las.ra as alpha, las.dec as delta, las.yAperMag3, las.j_1AperMag3, las.hAperMag3, las.kAperMag3, mj.mjdObs
FROM
```

```
    lasSource AS las, Multiframe AS mj, lasMergeLog AS l
```

```
WHERE j_1mfID=mj.multiframeID
```

```
    AND las.frameSetID=l.frameSetID
```

```
    AND yClassStat BETWEEN -3.0 AND 3.0
```

```
    AND j_1ClassStat BETWEEN -3.0 AND 3.0
```

```
    AND hClassStat BETWEEN -3.0 AND 3.0
```

```
    AND kClassStat BETWEEN -3.0 AND 3.0
```

```
    AND j_1Xi BETWEEN -1.0 AND +1.0
```

```
    AND hXi BETWEEN -1.0 AND +1.0
```

```
    AND kXi BETWEEN -1.0 AND +1.0
```

```
    AND j_1Eta BETWEEN -1.0 AND +1.0
```

```
    AND hEta BETWEEN -1.0 AND +1.0
```

```
    AND kEta BETWEEN -1.0 AND +1.0
```

```
    AND las.sourceID NOT IN (SELECT masterObjID FROM lasSourceXtwomass_psc WHERE distancemins < 0.01)
```

```
    AND (las.priOrSec=0 OR las.priOrSec=las.frameSetID)
```

```
    AND (las.yAperMag3-las.j_1AperMag3)>0.3
```

```
    AND ((las.yAperMag3-las.j_1AperMag3)>0.7 OR (las.j_1AperMag3-las.hAperMag3)<0.0)
```

```
    AND las.j_1AperMag3<16.0
```

```
    AND las.j_1AperMag3>10.5
```

```
    AND las.yAperMag3>11.3
```

```
    AND las.hAperMag3>10.2
```

```
    AND las.kAperMag3>9.7
```

```
    AND mj.mjdObs>0.0
```

```
    AND las.j_1E11<0.3
```

```
    AND las.yE11<0.3
```

```
    AND las.hE11<0.3
```

```
    AND las.kE11<0.3
```

APPENDIX B: DATA TABLES

Table B1. Objects identified in our LAS sample. All photometry is in the UKIDSS system (Hewett et al. 2006). ¹ denotes objects whose astrometric solutions was calculated using local reference stars while ² denotes those calculated using only global error estimates. Citation key - a:Schneider et al. (2002) b: Luyten (1979b) c: Tinney (1993) d: Fan et al. (2000) e: Bouy et al. (2003) f: Lepine & Shara (2005) g: Cruz et al. (2003) h: Hawley et al. (2002) i: Delfosse et al. (1999) j: Burgasser et al. (2004) k: Luyten (1979a) l Cruz et al. (2007) m: Wilson et al. (2003) n: Reid et al. (2008), p: Knapp et al. (2004), q: Gizis et al. (2000), r:Kirkpatrick et al. (2000).

Name	Position	μ "/yr	P.A. °	σ_{μ} "/yr	σ_{PA}	<i>Y</i>	<i>J</i>	<i>H</i>	<i>K</i>	note
ULAS2MASS1504+0923	15 04 10.17 +09 23 23.9	0.271	289.7	0.028 ¹	6.7	13.738	13.028	12.518	12.077	b
ULAS2MASS1505+1029	15 05 56.42 +10 29 40.0	0.106	183.4	0.016 ¹	13.5	15.387	14.682	14.119	13.707	
ULAS2MASS1510-0241	15 10 16.67 -02 41 07.7	0.399	272.2	0.014 ²	1.8	13.494	12.543	11.981	11.315	c
ULAS2MASS1511-0215	15 11 53.59 -02 15 31.8	0.145	105.8	0.017 ¹	8.2	15.629	14.923	14.414	13.986	
ULAS2MASS1513+0844	15 13 58.26 +08 44 34.4	0.136	157.1	0.018 ¹	7.9	15.624	14.753	14.186	13.762	
ULAS2MASS1514+1201	15 14 13.77 +12 01 45.0	0.161	260.4	0.018 ¹	6.6	15.795	14.803	14.230	13.732	
ULAS2MASS1516+0942	15 16 39.94 +09 42 11.6	0.181	286.7	0.020 ¹	4.7	14.626	13.696	13.170	12.689	
ULAS2MASS1518+0849	15 18 35.32 +08 49 08.0	0.187	139.2	0.027 ¹	8.1	16.157	15.364	14.905	14.434	
ULAS2MASS1518+0807	15 18 49.12 +08 07 43.5	0.175	216.5	0.021 ¹	6.5	15.831	15.129	14.649	14.247	
ULAS2MASS1520+0407	15 20 46.08 +04 07 52.3	0.097	195.6	0.013 ¹	10.2	14.418	13.635	13.070	12.676	
ULAS2MASS1521+0653	15 21 05.19 +06 53 07.0	0.120	277.8	0.022 ¹	9.7	15.955	15.177	14.569	14.154	
ULAS2MASS1524+1019	15 24 26.65 +10 19 49.4	0.114	308.5	0.016 ¹	8.1	14.771	14.015	13.473	13.027	
ULAS2MASS1524+0934	15 24 32.04 +09 34 37.6	0.155	169.0	0.022 ¹	5.5	15.762	14.948	14.442	14.048	
ULAS2MASS1526+0745	15 26 54.56 +07 45 41.2	0.174	206.0	0.021 ¹	6.2	16.006	15.139	14.496	14.022	
ULAS2MASS1529+0637	15 29 58.14 +06 37 00.3	0.194	225.7	0.030 ¹	8.9	16.212	15.500	15.046	14.629	
ULAS2MASS1531+0644	15 31 46.98 +06 44 23.1	0.164	123.2	0.023 ¹	7.4	14.847	14.142	13.604	13.218	f
ULAS2MASS1531+0915	15 31 54.73 +09 15 59.9	0.128	256.5	0.023 ¹	11.0	16.407	15.693	15.192	14.772	
ULAS2MASS1532+0837	15 32 17.20 +08 37 12.5	0.175	184.6	0.019 ¹	6.3	16.208	15.368	14.730	14.250	
ULAS2MASS1535+0932	15 35 41.48 +09 32 02.4	0.109	259.7	0.018 ¹	22.4	16.235	15.414	14.896	14.451	
ULAS2MASS1537+0604	15 37 05.80 +06 04 04.2	0.148	199.5	0.022 ¹	8.1	15.101	14.268	13.650	13.183	
ULAS2MASS1541+0620	15 41 43.77 +06 20 43.7	0.141	261.1	0.020 ¹	10.5	16.684	15.733	15.188	14.712	
ULAS2MASS1543+0939	15 43 15.91 +09 39 35.5	0.165	261.8	0.013 ¹	5.7	15.665	14.939	14.406	13.971	
ULAS2MASS1545-0124	15 45 13.52 -01 24 48.6	0.218	102.8	0.015 ¹	3.4	14.858	14.032	13.540	13.156	
ULAS2MASS1546+0317	15 46 25.93 +03 17 54.8	0.113	277.7	0.021 ¹	11.9	13.951	13.144	12.563	12.096	n
ULAS2MASS1547+1023	15 47 43.51 +10 23 31.8	0.156	232.4	0.015 ¹	6.0	15.710	14.971	14.477	14.044	
ULAS2MASS1553+0749	15 53 08.58 +07 49 30.6	0.123	303.0	0.021 ¹	9.7	14.908	14.054	13.491	13.050	
ULAS2MASS2143-0040	21 43 41.24 -00 40 44.4	0.123	54.5	0.017 ¹	10.5	14.598	13.883	13.369	12.998	
ULAS2MASS2147-0029	21 47 21.03 -00 29 46.3	0.186	26.2	0.027 ¹	10.4	15.274	14.535	13.942	13.529	
ULAS2MASS2148+0020	21 48 30.90 +00 20 53.8	0.177	140.8	0.020 ¹	6.7	16.314	15.397	14.748	14.244	
ULAS2MASS2203+0024	22 03 23.83 +00 24 05.5	0.163	235.8	0.012 ¹	4.5	15.738	14.979	14.467	14.051	c
ULAS2MASS2206-0037	22 06 09.69 -00 37 27.9	0.188	186.2	0.009 ¹	4.2	15.254	14.542	14.015	13.616	c
ULAS2MASS2208-0003	22 08 17.74 -00 03 09.2	0.126	89.0	0.022 ¹	7.6	15.399	14.695	14.122	13.742	
ULAS2MASS2214+0052	22 14 59.79 +00 52 34.4	0.241	96.1	0.021 ¹	3.0	14.731	14.000	13.479	13.106	c
ULAS2MASS2221-0022	22 21 16.88 -00 22 16.8	0.128	116.0	0.022 ¹	13.7	15.687	14.854	14.256	13.823	c
ULAS2MASS2237-0039	22 37 39.94 -00 39 50.8	0.122	129.8	0.017 ¹	7.7	14.713	13.913	13.357	12.933	c
ULAS2MASS2249+0025	22 49 35.39 +00 25 58.2	0.187	116.5	0.016 ¹	4.8	15.221	14.513	13.908	13.532	c
ULAS2MASS2258+0113	22 58 54.07 +01 13 49.9	0.252	164.3	0.013 ¹	3.2	14.695	13.875	13.365	12.883	
ULAS2MASS2300+0103	23 00 52.47 +01 03 13.8	0.271	121.1	0.028 ¹	5.0	15.704	14.838	14.228	13.733	
ULAS2MASS2304+0749	23 04 25.97 +07 49 00.9	0.237	96.7	0.022 ¹	5.0	15.414	14.464	13.817	13.330	
ULAS2MASS2308+0035	23 08 39.85 +00 35 34.3	0.218	123.0	0.035 ¹	10.0	16.399	15.550	14.939	14.497	
ULAS2MASS2314+0623	23 14 09.89 +06 23 09.4	0.161	90.7	0.031 ¹	10.4	16.428	15.613	15.027	14.542	
ULAS2MASS2316+1339	23 16 37.93 +13 39 15.3	0.091	241.3	0.013 ¹	8.1	15.895	15.185	14.590	14.179	
ULAS2MASS2318+1305	23 18 18.42 +13 05 03.8	0.142	217.6	0.016 ¹	6.6	15.212	14.471	13.937	13.560	
ULAS2MASS2321+0815	23 21 30.10 +08 15 45.2	0.137	88.8	0.025 ¹	11.4	16.037	15.249	14.706	14.289	
ULAS2MASS2331-0005	23 31 01.72 -00 05 30.1	0.251	98.3	0.012 ¹	4.3	15.728	14.998	14.496	14.082	
ULAS2MASS2331+1552	23 31 29.29 +15 52 21.9	0.158	234.8	0.012 ¹	4.1	16.034	14.976	14.303	13.736	
ULAS2MASS2332-0050	23 32 24.41 -00 50 25.2	0.098	86.5	0.013 ²	6.6	14.330	13.555	13.013	12.579	
ULAS2MASS2333+0050	23 33 58.48 +00 50 12.2	0.168	76.7	0.015 ¹	4.6	15.879	14.934	14.359	13.851	
ULAS2MASS2338+1604	23 38 24.77 +16 04 58.6	0.193	31.9	0.016 ¹	5.2	15.677	14.936	14.402	13.978	
ULAS2MASS2345+0055	23 45 39.08 +00 55 13.4	0.115	112.1	0.014 ²	6.3	14.632	13.677	13.087	12.523	n
ULAS2MASS2350+1441	23 50 49.62 +14 41 37.2	0.107	56.4	0.012 ¹	7.4	15.137	14.381	13.809	13.383	
ULAS2MASS2353+1511	23 53 38.48 +15 11 49.4	0.260	102.6	0.019 ¹	3.4	15.434	14.680	14.137	13.717	
ULAS2MASS2355+0115	23 55 42.54 +01 15 22.2	0.173	61.9	0.028 ¹	8.1	15.025	14.251	13.723	13.374	
ULAS2MASS2355+0016	23 55 46.28 +00 16 27.9	0.136	180.5	0.024 ¹	7.6	15.045	14.276	13.746	13.344	
ULAS2MASS2356-0034	23 56 17.60 -00 34 33.8	0.120	216.8	0.019 ¹	9.6	16.188	15.473	14.927	14.520	
ULAS2MASS2356-0058	23 56 34.70 -00 58 16.3	0.142	223.6	0.023 ¹	9.3	16.468	15.710	15.204	14.825	

Table B2. Objects identified in our GCS sample. All photometry is in the UKIDSS system (Hewett et al. 2006). ¹ denotes objects whose astrometric solutions was calculated using local reference stars while ² denotes those calculated using only global error estimates. * denotes the object has been identified as a potential cluster member. a: Stauffer et al. (1999) b: Barrado et al. (2002) c: Cossburn et al. (1997) d: Kirkpatrick et al. (1997) e: Ambartsumyan et al. (1973) f: Bouvier et al. (1998) g: Lepine & Shara (2005) h: Kirkpatrick et al. (2000) i: Luyten (1979b) j: Lepine, Shara & Rich (2002)

Name	Position	μ "/yr	P.A. °	σ_{μ} "/yr	σ_{PA}	<i>Y</i>	<i>J</i>	<i>H</i>	<i>K</i>	note
UGCS2MASS0319+5030	03 19 41.42 +50 30 43.8	0.194	147.1	0.008 ¹	2.3	16.109	15.210	14.532	14.013	a*
UGCS2MASS0319+4848	03 19 51.84 +48 48 22.0	0.220	247.6	0.008 ¹	2.2	14.748	13.899	13.368	12.906	b*
UGCS2MASS0328+4841	03 28 02.35 +48 41 05.8	0.309	130.7	0.013 ¹	2.4	15.629	14.896	14.415	13.980	b*
UGCS2MASS0335+4809	03 35 38.61 +48 09 20.7	0.110	142.0	0.014 ¹	7.4	14.569	13.829	13.263	12.836	
UGCS2MASS0336+2233	03 36 25.95 +22 33 17.2	0.243	132.7	0.007 ¹	1.7	15.808	14.991	14.471	13.977	
UGCS2MASS0339+2457	03 39 52.92 +24 57 27.0	0.178	107.6	0.028 ¹	4.2	13.622	12.786	12.212	11.744	d
UGCS2MASS0340+2215	03 40 43.38 +22 15 07.0	0.154	268.4	0.014 ²	4.1	13.819	13.089	12.545	12.142	e*
UGCS2MASS0341+3207	03 41 16.45 +32 07 55.2	0.137	214.2	0.008 ¹	3.6	15.905	15.137	14.594	14.093	
UGCS2MASS0342+2248	03 42 10.17 +22 48 44.6	0.089	166.9	0.015 ¹	10.7	16.355	15.543	14.932	14.469	
UGCS2MASS0343+3037	03 43 57.35 +30 37 19.0	0.083	155.5	0.006 ¹	6.5	15.612	14.763	14.117	13.586	
UGCS2MASS0345+2540	03 45 43.12 +25 40 23.1	0.109	248.5	0.020 ¹	6.0	15.009	13.924	13.240	12.663	c*
UGCS2MASS0347+3406	03 47 54.07 +34 06 47.7	0.096	153.7	0.010 ¹	5.9	16.262	15.304	14.701	14.139	
UGCS2MASS0350+3149	03 50 02.97 +31 49 58.8	0.124	62.5	0.009 ¹	4.1	16.255	15.490	14.972	14.516	
UGCS2MASS0351+2544	03 51 57.26 +25 44 15.0	0.373	109.0	0.024 ¹	3.3	15.045	14.343	13.928	13.504	
UGCS2MASS0352+2359	03 52 07.88 +23 59 13.1	0.091	165.4	0.016 ¹	5.7	15.371	14.636	14.066	13.640	f*
UGCS2MASS0354+2316	03 54 01.47 +23 16 33.5	0.206	107.0	0.009 ¹	5.1	13.895	13.046	12.452	11.980	d*
UGCS2MASS0401+2322	04 01 43.07 +23 22 12.4	0.112	126.8	0.010 ¹	4.8	14.151	13.413	12.869	12.442	
UGCS2MASS0402+2837	04 02 58.77 +28 37 53.5	0.099	115.2	0.013 ¹	5.3	15.789	15.069	14.549	14.114	
UGCS2MASS0403+2616	04 03 07.86 +26 16 08.0	0.654	135.4	0.007 ¹	0.7	14.207	13.489	13.068	12.708	j
UGCS2MASS0407+2734	04 07 48.22 +27 34 05.2	0.088	191.3	0.014 ¹	11.9	16.183	15.453	14.918	14.509	
UGCS2MASS0408+2447	04 08 40.01 +24 47 32.3	0.081	139.1	0.008 ¹	5.3	16.275	15.499	14.592	14.199	
UGCS2MASS0409+2104	04 09 09.57 +21 04 37.9	0.184	150.6	0.007 ¹	2.4	16.622	15.376	14.502	13.781	h
UGCS2MASS0409+2446	04 09 33.19 +24 46 33.7	0.144	146.3	0.011 ¹	4.5	16.268	15.468	14.878	14.376	
UGCS2MASS0427+2900	04 27 23.86 +29 00 27.0	0.102	142.3	0.011 ¹	6.5	16.296	15.520	14.862	14.448	
UGCS2MASS0433+2933	04 33 37.21 +29 33 13.1	0.120	180.6	0.010 ¹	2.7	14.395	13.608	12.651	12.317	
UGCS2MASS0434+2937	04 34 54.42 +29 37 22.2	0.239	137.8	0.009 ¹	2.2	16.325	15.595	14.911	14.522	
UGCS2MASS0435+2956	04 35 23.50 +29 56 27.4	0.084	154.1	0.010 ¹	5.5	15.415	14.674	13.942	13.522	
UGCS2MASS0436+3019	04 36 26.30 +30 19 56.3	0.115	129.5	0.010 ¹	5.1	16.030	15.312	14.840	14.433	
UGCS2MASS0535-0551	05 35 06.67 -05 51 09.8	0.159	173.7	0.016 ¹	5.8	15.599	14.884	14.344	13.920	
UGCS2MASS0536-0454	05 36 16.38 -04 54 34.2	0.083	101.3	0.013 ¹	9.8	16.418	15.623	14.714	14.318	
UGCS2MASS0536-0328	05 36 37.18 -03 28 50.8	0.354	157.0	0.017 ¹	2.2	15.053	14.186	13.554	13.087	
UGCS2MASS0537-0403	05 37 47.55 -04 03 26.9	0.157	184.2	0.012 ¹	5.9	15.300	14.564	13.972	13.564	
UGCS2MASS0541-0305	05 41 55.81 -03 05 14.2	0.168	324.2	0.009 ¹	3.1	16.387	15.381	14.271	13.784	
UGCS2MASS0830+1821	08 30 29.92 +18 21 22.8	0.118	124.4	0.018 ¹	7.7	15.751	14.821	14.185	13.693	
UGCS2MASS0831+2036	08 31 28.42 +20 36 54.1	0.097	177.5	0.010 ¹	6.9	16.136	15.365	14.900	14.477	
UGCS2MASS0835+2030	08 35 29.18 +20 30 49.0	0.111	193.8	0.018 ¹	7.0	14.555	13.829	13.262	12.871	
UGCS2MASS0835+2224	08 35 45.28 +22 24 30.9	0.159	261.5	0.015 ¹	5.8	16.490	15.559	14.927	14.443	
UGCS2MASS0838+1852	08 38 08.04 +18 52 04.9	0.107	184.1	0.017 ¹	7.5	15.620	14.814	14.259	13.844	
UGCS2MASS0840+2024	08 40 48.29 +20 24 10.4	0.411	121.8	0.012 ¹	2.0	16.411	15.661	15.248	14.871	
UGCS2MASS0848+2150	08 48 16.74 +21 50 51.0	0.339	152.1	0.019 ¹	3.3	14.458	13.738	13.229	12.850	g
UGCS2MASS1548-2148	15 48 45.71 -21 48 08.8	0.156	199.7	0.017 ¹	7.4	16.491	15.786	15.256	14.849	
UGCS2MASS1550-2201	15 50 11.50 -22 01 21.9	0.082	194.1	0.016 ¹	13.9	16.337	15.548	14.980	14.573	
UGCS2MASS1620-2847	16 20 03.77 -28 47 21.5	0.100	221.6	0.010 ¹	5.8	14.852	14.088	13.553	13.130	
UGCS2MASS1625-2117	16 25 46.88 -21 17 26.3	0.112	164.6	0.017 ¹	7.4	16.253	15.507	14.888	14.431	
UGCS2MASS1625-2400	16 25 50.24 -24 00 08.5	0.178	261.1	0.016 ²	3.6	12.799	11.892	10.973	10.545	i
UGCS2MASS1626-2352	16 26 25.13 -23 52 37.8	0.133	100.4	0.021 ²	6.5	12.585	11.802	11.072	10.704	
UGCS2MASS1628-2116	16 28 56.61 -21 16 09.9	0.097	152.5	0.017 ¹	9.1	14.650	13.935	13.168	12.747	
UGCS2MASS1630-2120	16 30 17.69 -21 20 01.4	0.153	259.8	0.013 ¹	3.8	15.695	14.524	13.780	13.174	
UGCS2MASS1631-2404	16 31 46.42 -24 04 47.5	0.104	87.5	0.013 ¹	5.9	16.530	15.735	14.859	14.396	
UGCS2MASS1632-2358	16 32 21.35 -23 58 23.8	0.160	197.2	0.012 ¹	4.8	16.390	15.554	14.993	14.494	
UGCS2MASS1637-2200	16 37 53.00 -22 00 14.7	0.084	252.1	0.011 ¹	9.4	15.997	15.112	14.259	13.821	
UGCS2MASS1639-2502	16 39 29.17 -25 02 12.3	0.093	142.9	0.014 ¹	7.5	16.189	15.466	14.800	14.489	
UGCS2MASS1640-2442	16 40 22.44 -24 42 48.5	0.083	209.8	0.014 ¹	10.8	16.425	15.599	14.930	14.441	
UGCS2MASS1751+0504	17 51 00.24 +05 04 23.0	0.173	204.4	0.029 ¹	9.1	16.752	15.733	15.073	14.503	